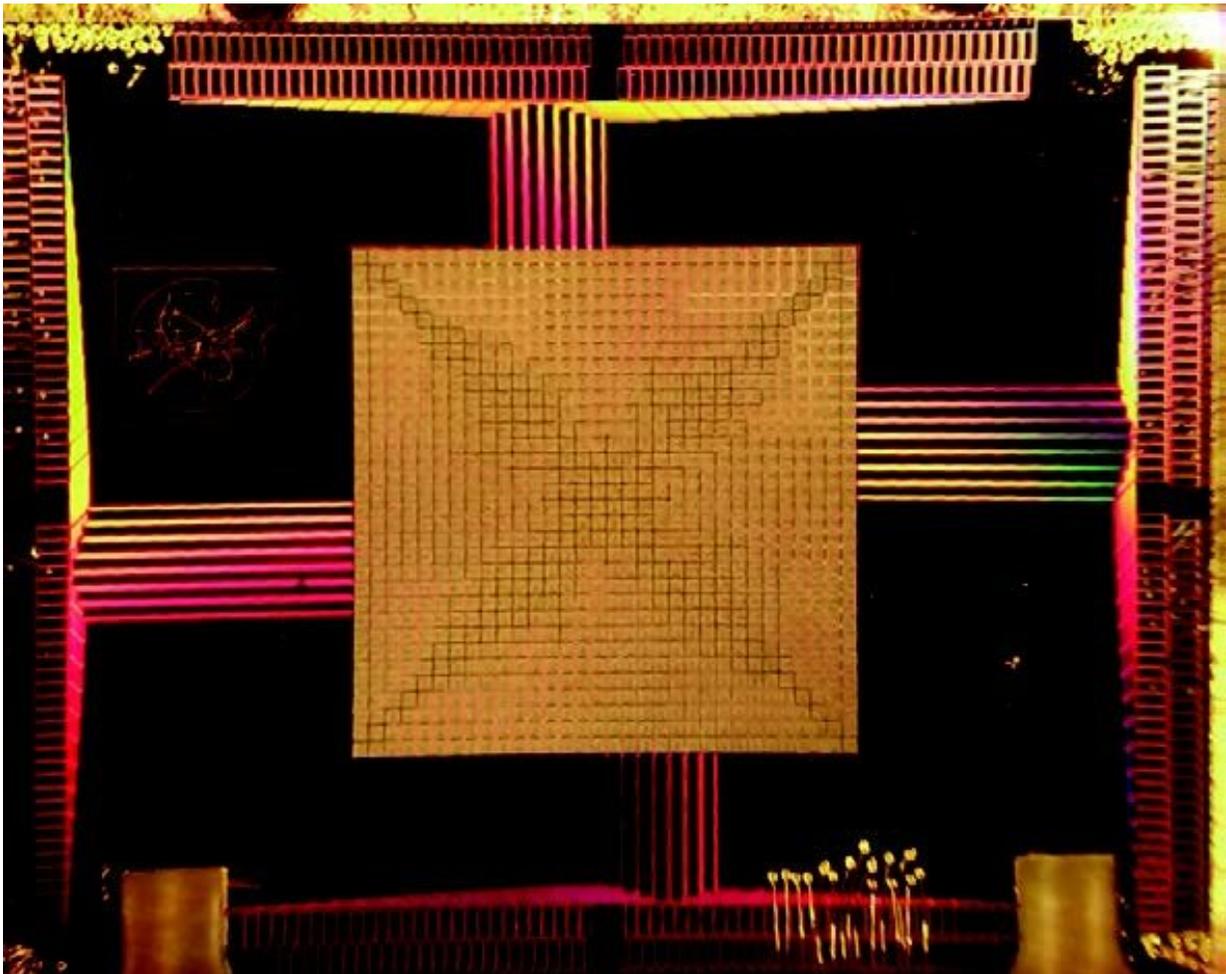


# An X-ray camera that can resolve tens of thousands of X-ray colors

February 13 2018

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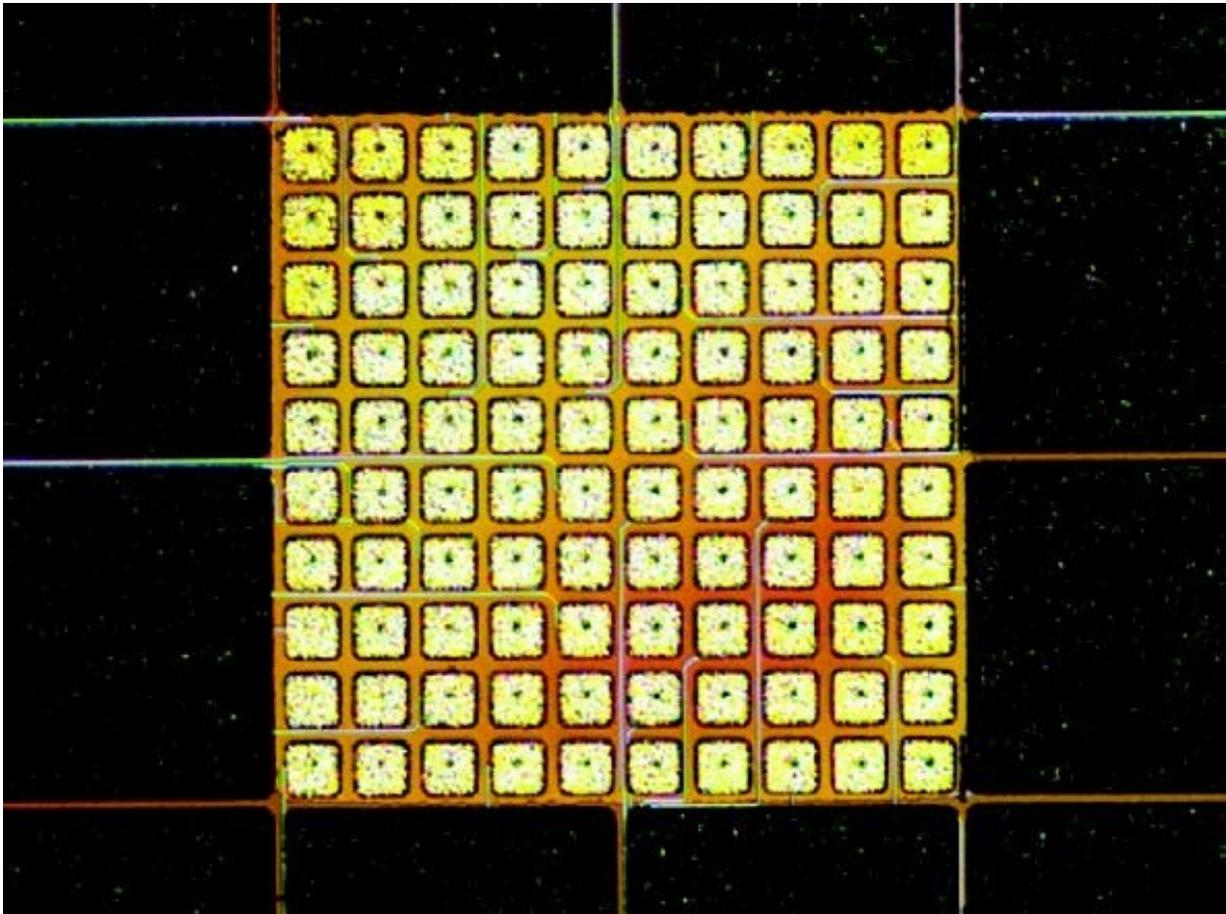


A 32x32 TES X-ray microcalorimeter array that is a candidate for use in the Athena X-IFU demonstration model. The individual, closely packed Au/Bi absorbers are arrayed on a 0.25 mm grid. Credit: NASA

NASA is part of an international team developing a cutting edge microcalorimeter X-ray camera that will provide extraordinarily detailed information about energetic cosmic phenomena.

An X-ray microcalorimeter is a non-dispersive spectrometer that uses an equilibrium approach to energy measurement—the energy of an X-ray photon heats an isolated thermal mass, and the temperature change is measured. The ultimate energy resolution is determined by how well the temperature pulse can be measured against a background of thermal fluctuations; thus, high-resolution spectrometers must be operated at very low temperatures (

With each improvement, new mission concepts are developed that require even larger arrays. The NASA/Japan Aerospace Exploration Agency (JAXA) Soft X-ray Spectrometer (SXS) instrument onboard the JAXA Hitomi mission had 36 pixels, but the X-ray Integral Field Unit (X-IFU) instrument that will fly on the European Space Agency's Athena mission requires an array of about 4000 pixels, each about 0.25 mm wide (covering 5 arc seconds of the sky). X-IFU will be a groundbreaking X-ray camera capable of distinguishing tens of thousands of X-ray colors. As part of the X-IFU consortium, NASA is developing the superconducting transition edge sensor (TES) array employed on the instrument. These sensors, composed of Molybdenum/Gold TES thermometers and Gold/Bismuth X-ray absorbers, achieve better than 2.5 eV resolution.

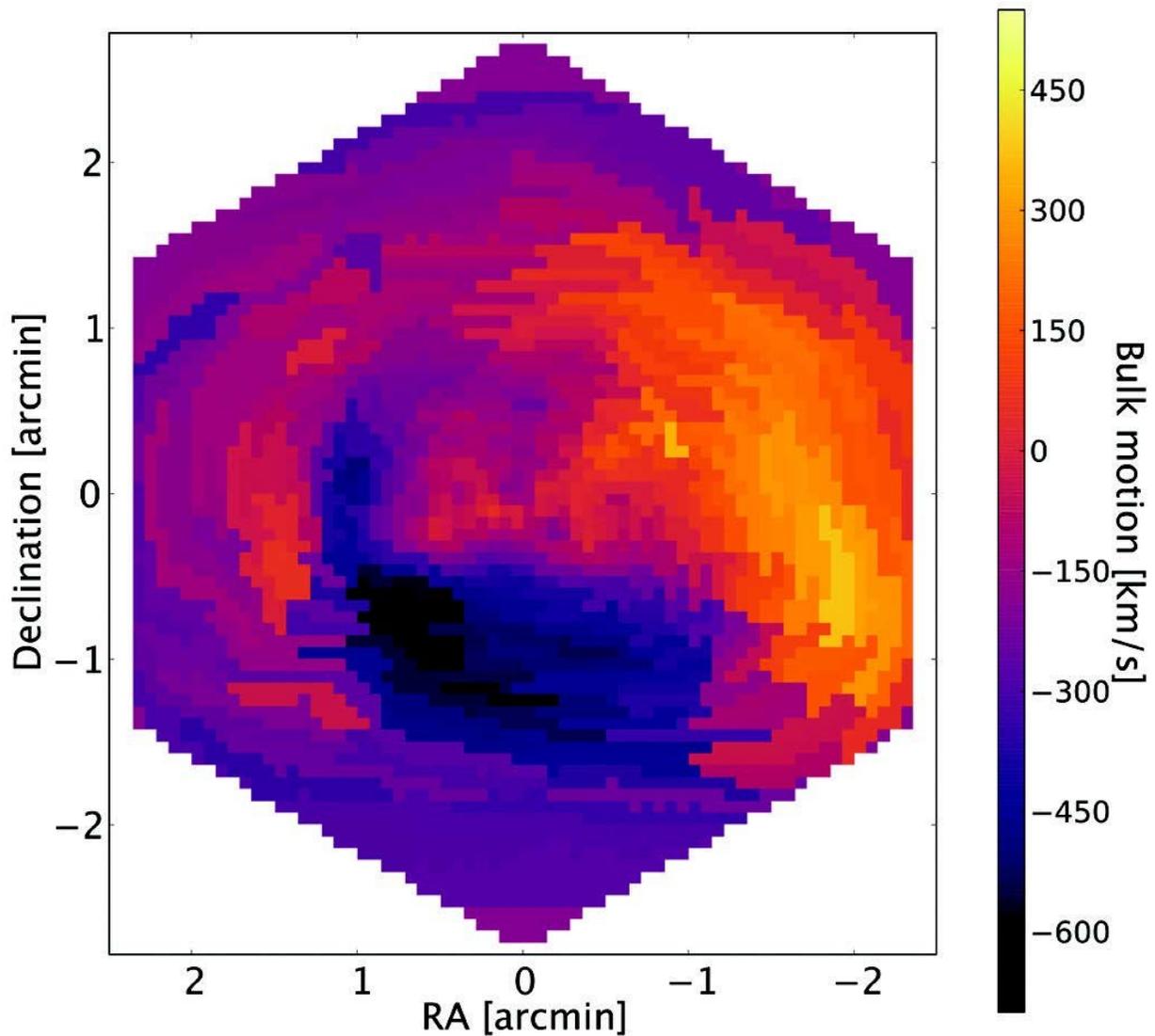


An example of a hybrid array consisting of a 10 x 10 array of small-pixel TESs with Au absorbers (0.05 mm pitch) surrounded by an array of larger pixels with composite Au/Bi absorbers (0.25 mm pitch). Such an array was considered for X-IFU to provide a central region optimized for higher spectral resolution or higher count rates. Credit: NASA

Cosmic phenomena that produce X rays characterize the evolution of cosmic structures on both large and small scales. High-resolution X-ray spectroscopy can determine density and temperature, identify ions and determine their velocities, and enable scientists to study effects such as turbulence or the environment near supermassive black holes. Combining imaging with spectroscopy, a microcalorimeter instrument

probes dynamics and variations within spatially extended objects such as supernova remnants and galaxy clusters with unprecedented sensitivity.

In 2016, the NASA team focused on working with partners at SRON, the Netherlands Institute for Space Research, to prepare for an X-IFU demonstration model incorporating a kilopixel TES array. Because the planned readout for X-IFU uses frequency-division multiplexing, which involves applying alternating voltages to the TES thermometers, the near-term focus has been on determining the optimal pixel design for that mode of operation. Important progress was also achieved using backup multiplexing technologies that apply a constant voltage to the TES thermometers (time-division and code-division). A time-division multiplexing demonstration of a column of 32 TES pixels achieved 2.55 eV average energy resolution at 6 keV at a speed appropriate for the original X-IFU baseline. The team completed the layout for a full-sized X-IFU prototype array, and in the next year these prototypes will be fabricated and tested. The team also successfully demonstrated that pixels with different characteristics (width, X-ray absorber materials and thickness, and superconducting transition temperature) can be incorporated into a single array, should it be determined to be optimal to do so on X-IFU or another mission.



Simulated map of line-of-sight velocities in the X-ray emitting gas in a galaxy cluster like the Perseus cluster, as determined from the high-resolution X-ray spectrum that the X-IFU will measure for each pixel. (Credit: left panel of Fig 2 in Barret et al. 2016, Proc. SPIE. 9905, Space Telescopes and Instrumentation 2016: Ultraviolet to Gamma Ray, 99052F). Credit: NASA

Provided by NASA

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